

First-year Sloan Digital Sky Survey-II (SDSS-II) supernova results: consistency and constraints with other intermediate-redshift datasets

H. Lampeitl^{1*}, R.C. Nichol¹, H.-J. Seo², T. Giannantonio^{1,3}, C. Shapiro¹,
B. Bassett^{4,5}, W.J. Percival¹, T.M. Davis^{6,7}, B. Dilday⁸, J. Frieman^{2,9},
P. Garnavich¹⁰, M. Sako¹¹, M. Smith^{1,4}, J. Sollerman^{7,12}, A.C. Becker¹³,
D. Cinabro¹⁴, A.V. Filippenko¹⁵, R.J. Foley^{15,16,17}, C.J. Hogan², J.A. Holtzman¹⁸,
S.W. Jha⁸, K. Konishi^{19,20}, J. Marriner², M.W. Richmond²¹, A.G. Riess^{22,23},
D.P. Schneider²⁴, M. Stritzinger^{7,25}, K.J. van der Heyden²⁶, J.T. VanderPlas¹³,
J.C. Wheeler²⁷, C. Zheng²⁸

¹*Institute of Cosmology and Gravitation, University of Portsmouth, Portsmouth, PO1 3FX, UK*

²*Center for Particle Astrophysics, Fermi National Accelerator Laboratory, P.O. Box 500, Batavia, IL 60510, USA*

³*Argelander-Institut für Astronomie, Universität Bonn, Auf dem Hügel 71, D-53121 Bonn, Germany*

⁴*Department of Mathematics and Applied Mathematics, University of Cape Town, Rondebosch 7701, South Africa*

⁵*South African Astronomical Observatory, P.O. Box 9, Observatory 7935, South Africa*

⁶*School of Mathematics and Physics, University of Queensland, Brisbane QLD 4072, Australia*

⁷*Dark Cosmology Centre, Niels Bohr Institute, University of Copenhagen, DK-2100, Copenhagen, Denmark*

⁸*Department of Physics and Astronomy, Rutgers the State University of New Jersey, 136 Frelinghuysen Road, Piscataway, NJ 08854, USA*

⁹*Kavli Institute for Cosmological Physics, The University of Chicago, 5640 South Ellis Avenue Chicago, IL 60637, USA*

¹⁰*Physics Department, University of Notre Dame, Notre Dame, IN 46556, USA*

¹¹*University of Pennsylvania, Department of Physics and Astronomy, 209 South 33rd Street, Philadelphia, PA 19096, USA*

¹²*Department of Astronomy, AlbaNova, Stockholm University, SE-106 91 Stockholm, Sweden*

¹³*Department of Astronomy, University of Washington, Box 351580, Seattle, WA 98195, USA*

¹⁴*Wayne State University, Department of Physics and Astronomy, Detroit, MI 48202, USA*

¹⁵*Department of Astronomy, University of California, Berkeley, CA 94720-3411, USA*

¹⁶*Harvard-Smithsonian Center for Astrophysics, 60 Garden Street, Cambridge, MA 02138*

¹⁷*Clay Fellow*

¹⁸*Dept. of Astronomy, MSC 4500, New Mexico State University, P.O. Box 30001, Las Cruces, NM 88003, USA*

¹⁹*Institute for Cosmic Ray Research, The University of Tokyo, 5-1-5 Kashiwa, Kashiwa City, Chiba 277-8582, Japan*

²⁰*Department of Physics, Graduate School of Science, The University of Tokyo, 7-3-1 Hongo, Bunkyo, Tokyo 113-0033, Japan*

²¹*Physics Department, Rochester Institute of Technology, Rochester, NY 14623, USA*

²²*Department of Physics and Astronomy, Johns Hopkins University, Baltimore, MD 21218, USA*

²³*Space Telescope Science Institute, 3700 San Martin Drive, Baltimore, MD 21218, USA*

²⁴*Department of Astronomy and Astrophysics, 525 Davey Laboratory, Pennsylvania State University, University Park, PA 16802, USA*

²⁵*Las Campanas Observatory, Carnegie Observatories, Casilla 601, La Serena, Chile*

²⁶*Astronomy Department, University of Cape Town, Private Bag X3, Rondebosch 7701, South Africa*

²⁷*Department of Astronomy, McDonald Observatory, University of Texas, Austin, TX 78712, USA*

²⁸*Kavli Institute for Particle Astrophysics and Cosmology, Stanford University, CA 94305-4060*

ABSTRACT

We present an analysis of the luminosity distances of Type Ia Supernovae from the Sloan Digital Sky Survey-II (SDSS-II) Supernova Survey in conjunction with other intermediate redshift ($z < 0.4$) cosmological measurements including redshift-space distortions from the Two-degree Field Galaxy Redshift Survey (2dFGRS), the Integrated Sachs-Wolfe (ISW) effect seen by the SDSS, and the latest Baryon Acoustic Oscillation (BAO) distance scale from both the SDSS and 2dFGRS. We have analysed the SDSS-II SN data alone using a variety of “model-independent” methods and find evidence for an accelerating universe at $>97\%$ level from this single dataset. We find good agreement between the supernova and BAO distance measurements, both consistent with a Λ -dominated CDM cosmology, as demonstrated through an analysis of the distance duality relationship between the luminosity (d_L) and angular diameter (d_A) distance measures. We then use these data to estimate w within this restricted redshift range ($z < 0.4$). Our most stringent result comes from the combination of all our intermediate-redshift data (SDSS-II SNe, BAO, ISW and redshift-space distortions), giving $w = -0.81_{-0.18}^{+0.16}(\text{stat}) \pm 0.15(\text{sys})$ and $\Omega_M = 0.22_{-0.08}^{+0.09}$ assuming a flat universe. This value of w , and associated errors, only change slightly if curvature is allowed to vary, consistent with constraints from the Cosmic Microwave Background. We also consider more limited combinations of the geometrical (SN, BAO) and dynamical (ISW, redshift-space distortions) probes.

Key words: (stars:) supernovae: general, (cosmology:) observations, distance-scale, cosmological parameters, large-scale structure of universe

1 INTRODUCTION

It is now widely believed that the late-time expansion of the universe is accelerating. General Relativity (GR) implies that the acceleration is driven by “dark energy” (DE) — an unknown energy component in the universe with a negative effective pressure. Describing dark energy by an equation-of-state parameter of $w(z) = p/(\rho c^2)$ requires that $w < -1/3$. Alternatively, accelerated expansion could be an indication that GR is not the correct theory of gravity or that we have applied GR incorrectly in a cosmological context (see recent reviews of DE by Peebles & Ratra 2003, Uzan 2006, Copeland et al. 2006, and Frieman, Turner, & Huterer 2008).

Over the last decade, the most direct way of studying this acceleration of the expansion of the universe, and therefore DE, has been using Type Ia Supernovae (SNe), as they have been shown by many authors to be well-calibrated “standard candles” in the universe, i.e., their relative distances can be determined from the dependence of their peak luminosity on the shape of the light curve. This method was used to great effect by astronomers in 1998 to provide the first evidence for an accelerated universe (Riess et al. 1998; Perlmutter et al. 1999; see Filippenko 2005 for a review).

Briefly, a Type Ia supernova occurs when a white dwarf star in a close binary system accretes enough mass from its companion to undergo a thermonuclear explosion in the core. Both Phillips (1993) and Hamuy et al. (1993) have shown that such explosions can serve as consistent light sources in the universe to high accuracy. This is achieved by transforming the measured light curve of the explosion into the rest frame of the supernova (so called K-corrections) and

then correcting the luminosity at maximum as a function of the shape of the rest-frame light curve.

Several techniques now exist for fitting SN light curves known under different acronyms (Δm_{15} , Hamuy et al. 1996; MLCS, Riess, Press, & Kirshner 1996; stretch, Perlmutter et al. 1997; CMAGIC, Wang et al. 2003; BATM, Tonry et al. 2003; SALT, Guy et al. 2005; ΔC_{12} , Wang et al. 2006; SALT2, Guy et al. 2007; SiFTO, Conley et al. 2008). In this analysis, we consider MLCS2k2 (Jha, Riess, & Kirshner 2007; Riess, Press, & Kirshner 1996), which is among the most commonly used, and best tested, of these various techniques.

Recently, several dedicated SN surveys have been carried out to confirm and extend the earlier detections of an accelerating universe (HST, Riess et al. 2004, 2007; SNLS, Astier et al. 2006; ESSENCE, Wood-Vasey et al. 2007) as well as new compilations of existing SN datasets (Davis et al. 2007; Kowalski et al. 2008; Hicken et al. 2009). In addition to supernovae, observations of Baryon Acoustic Oscillations (BAO) can be used to measure distances in the universe (Blake & Glazebrook 2003; Seo & Eisenstein 2003; Hu & Haiman 2003). The BAO are caused by sound waves in the early universe which leave a preferred scale in the distribution of matter equal to the sound horizon at recombination. Today, this scale corresponds to $\sim 100/h$ Mpc (Hubble constant at present: $H_0 = 100h$ km/s/Mpc) and can thus be used as a “standard ruler” throughout the universe. The BAO signature has been detected in the clustering of galaxy clusters by Miller, Nichol, & Batuski (2001), in the correlation of galaxies in the Sloan Digital Sky Survey (SDSS, York et al. 2000) by Eisenstein et al. (2005), Hütsi (2006), Padmanabhan et al. (2007), and Blake et al. (2007), and in the Two-degree Field Galaxy Redshift Survey (2dFGRS, Colless et al. 2001) by Cole et al. (2005).

In addition to the geometrical methods discussed above,

* E-mail: Hubert.Lampeitl@port.ac.uk

observations of the dynamical properties of matter can provide constraints on the matter density of the universe and w (assuming General Relativity is the appropriate theory of gravity). For example, the growth rate of structure in the universe can be observed via the coherent infall of galaxies into large clusters and superclusters of galaxies seen in redshift surveys (Kaiser 1987). Also, the growth of structure can be measured using the late-time Integrated Sachs-Wolfe (ISW) effect (Sachs & Wolfe 1967), which has now been detected to high significance through the cross-correlation of galaxy surveys with the Cosmic Microwave Background (CMB) (see Giannantonio et al. 2008 and references therein). The ISW is sensitive to deviations from a matter-dominated, Einstein-de Sitter universe ($\Omega_M = 1$, where Ω_M is the matter density at present divided by the critical density.).

Taken together, the present combination of cosmological measurements suggests we live in a flat universe, dominated by a cosmological constant (Λ) with the energy density in matter and Λ known to a statistical accuracy of better than a few percent (see Dunkley et al. 2009). However, several of these cosmological measurements, especially SNe, are becoming limited by their systematic uncertainties which are now dominating, e.g., Hicken et al. (2009) showed that the best combination of available SNe and BAO measurements provide $1 + w = 0.013^{+0.066}_{-0.068}$ but with a systematic uncertainty of 0.11. Therefore, it is clear that future cosmological surveys must resolve these systematic errors through new observations and better analysis methods to mitigate their effect.

This paper is one of three complementary papers focused on the cosmological analysis of a new sample of intermediate supernova distances recently obtained by the SDSS-II Supernova Survey (see Section 2 for details). Our analysis differs from those presented in our companion papers of Kessler et al. (2009) and Sollerman et al. (2009), as we first study the cosmological information obtained solely from the SDSS-II SN sample, and then in combination with other cosmological probes over the same redshift range ($z < 0.4$). Alternatively, Kessler et al. (2009) presents a detailed examination of the impact of both statistical and systematic errors on deriving standard cosmological constraints based on the combination of the SDSS-II SN with most of the currently available high and low redshift SNe and which are all analysed in a consistent way. Sollerman et al. (2009) then use the same compilation of data to study an expanded set of exotic cosmological models, in combination with a wider variety of other cosmological information. Our approach is also complementary to the many other analyses in the literature (e.g. Davis et al. 2007, Kowalski et al. 2008, Hicken et al. 2009) that have used data from all possible sources.

In our approach we concentrate on the information from cosmological measurements that cover the same range of redshifts as the SDSS SN sample. Our aim is not to derive the most stringent limit on cosmological parameters available, but rather to verify that if we restrict ourselves to probes coming from a small and similar redshift slice the results on cosmological parameters remain stable and consistent. This approach is warranted because of the growing emphasis on controlling systematic uncertainties in the analysis of cosmological data. There are clearly a number of systematic uncertainties that could affect the use of SNe as cos-

mological probes which likely depend on, or change with, redshift, including SN evolution (e.g., changes in the metallicity of progenitor stars Timmes, Brown, & Truran (2003); Howell et al. (2009); Sullivan et al. (2009)), intergalactic dust (Conley et al. 2007; Holwerda 2008), Malmquist bias and the effects of gravitational lensing and peculiar velocities (Hui & Greene 2006). Moreover, the photometric uncertainties associated with combining SN data from multiple surveys, over a range of redshifts, is already seen as the main limitation in using presently available datasets (see Hicken et al. 2009). Our analysis addresses this issue by focusing exclusively on the SDSS SN sample, which is derived from a well-understood and stable photometric system. The SDSS has a relative photometric accuracy of better than 2% in griz, and 4% for the u-band (Padmanabhan et al. 2007), while the absolute calibration is also known to be of the order of 1%, leading to a homogeneous set of SN light-curves with high photometric accuracy (see Holtzman et al. 2008). This set of data is robust to uncertainties in light-curve fitting. For example, the MLCS2k2 and SALT2 fits to the SDSS-only sample are shown to agree well (see Section 10 in Kessler et al. 2009) which is not the case when the higher redshift SN samples are added.

The outline of this paper is as follows. In Section 2, we describe the SDSS-II SN data and use that data in Section 3 to study the cosmic acceleration in the Universe. Section 4 then compares the SDSS-II SN luminosity distances to the BAO distances from the SDSS and 2dFGRS, checking the distance duality relation. We then derive in Section 5 constraints on w by combining the best-fit luminosity distances for SDSS-II SNe with the growth rate of structure measurements taken from Hawkins et al. (2003) and a new measurement of the ISW effect taken from Giannantonio et al. (2008). We conclude in Section 6.

2 THE FIRST-YEAR SDSS SN DATA

The SDSS-II Supernova Survey (Frieman et al. 2008) was part of the SDSS-II project and was focused on constructing a large sample of intermediate-redshift SNe ($0.045 < z < 0.42$). One of the strengths of the SDSS-II SN Survey is that it builds upon the existing (and stable) infrastructure from the original SDSS (Fukugita et al. 1996; Gunn et al. 1998; York et al. 2000; Lupton et al. 2001; Hogg et al. 2001; Eisenstein et al. 2001; Strauss et al. 2002; Pier et al. 2003; Gunn et al. 2006). The SDSS-II SN Survey is based on repeat imaging of ‘‘Stripe 82’’, a region of the original SDSS with significantly deeper photometry than the regular SDSS survey. This region is $\sim 120^\circ$ long and 2.5° wide, centered along the celestial equator and extending from 20^{hrs} to 4^{hrs} in right ascension (passing through 0^{hrs}).

The SDSS-II SN Survey was carried out in three observing campaigns from September through December in each of 2005, 2006, and 2007 (there were also some observations for a short period in 2004). The new imaging data were initially reduced using the standard SDSS pipelines (Stoughton et al. 2002; Smith et al. 2002; Ivezić et al. 2004; Tucker et al. 2006; Padmanabhan et al. 2008; Adelman-McCarthy et al. 2008), followed by specific image-subtraction software to identify transient objects (Sako et al. 2008). To determine the nature of these transients, the data were both visu-

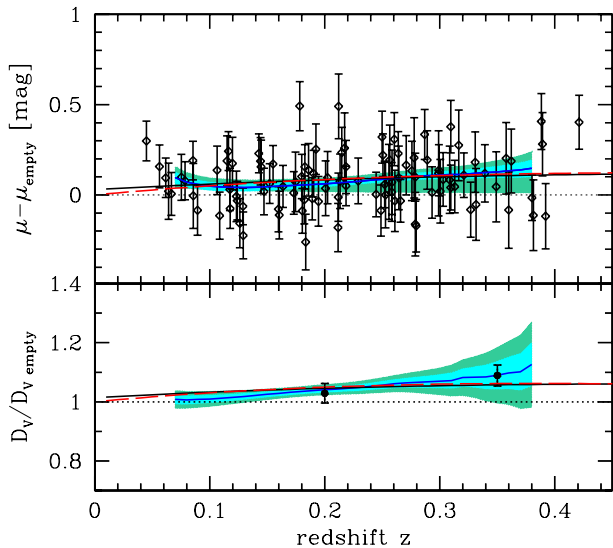


Figure 1. Residual Hubble diagram with respect to an empty universe for the 103 Type Ia SNe from the first year of operation of the SDSS-II SN Survey. The red line shows a Λ CDM model with $(\Omega_M, \Omega_\Lambda) = (0.3, 0.7)$, similar to our best-fit model given in Table 1. The blue line is the best fit to the data derived from the “sliding window” technique described in Section 4.1, and the cyan and green shaded regions correspond to the 1σ and 2σ confidence intervals, respectively. The black line indicates an expansion history with a deceleration parameter of $q_0 = -0.34$ as described in Section 3. The lower panel shows the same parameterization but now converted to D_V according to Eq. (9). The two data points represent the BAO measurements from Percival et al. (2009).

ally inspected and fit with SN models. Subsequently, objects with a high probability of being a SN Ia (Sako et al. 2008) were spectroscopically observed using a variety of telescopes around the world (Frieman et al. 2008; Zheng et al. 2008).

In this paper, we only consider the 2005 observing campaign (the first year of operation) as the data from other years is still being collated and analysed. That year, the SDSS-II SN Survey discovered 130 spectroscopically confirmed type Ia supernovae and an additional 16 spectroscopically probable SNe Ia. In Kessler et al. (2009) distance moduli are obtained for 103 of the spectroscopically confirmed SNe Ia that pass stringent light-curve quality cuts, using the MLCS2k2 light-curve fitting routine. In the upper panel of Fig. 1 we show for these 103 SNe the residuals of the distance modules with respect to an empty universe. We refer the reader to Section 9 of Kessler et al. (2009) for an extensive discussion of systematic effects caused by changing the various light-curve fitting parameters.

We also direct the reader to Section 11 of Kessler et al. (2009) for a comparison of the MLCS2k2 and SALT2 (Guy et al. 2007) light-curve fitters. They show that for the SDSS-only data the systematic difference between these two light-curve fitting methods is only 0.04 in w , while Fig. 42 of their paper shows the two methods give consistent distance moduli for the same SDSS-II supernovae. We also highlight that the two methods give similar contours in Fig. 26a and 35a of Kessler et al. (2009) when comparing the full cosmological fits to the SDSS-only data. This motivates the analysis in this paper and means our results are unaffected by the choice of light-curve fitter used. We restrict

the analysis to the MLCS2k2 reduction taken from Table 10 in Kessler et al. (2009), as these data include corrections for selection effects.

When using these SDSS-II SNe for cosmological fitting, we calculate the confidence intervals via the χ^2 statistic,

$$\chi^2 = \sum_{i=1}^N \frac{(\mu_{\text{LC},i}(z_i) - \mu_{\text{model}}(z_i, \vec{x}))^2}{\sigma_{\text{LC},i}^2 + \sigma_{\text{sp}}^2 + \sigma_{\text{int}}^2}, \quad (1)$$

where μ_{LC} and σ_{LC} are the distance moduli and errors, respectively, derived from the light-curve fitting method (see Kessler et al. 2009), and μ_{model} are the expected distance moduli according to parameters \vec{x} of the assumed cosmological model. Uncertainties in the measured spectroscopic redshift and peculiar velocity of the SN are taken into account using

$$\sigma_{\text{sp}} = \frac{5}{\ln(10)} \frac{(1+z)}{z(1+z/2)} \sqrt{(\Delta z)^2 + (\Delta v_p)^2/c^2}, \quad (2)$$

where Δz is the measurement uncertainty in redshift and Δv_p is the characteristic amplitude of the peculiar velocities, which we take to be 360 km s^{-1} (see, e.g., Masters et al. 2006). We investigated the effect of correlated peculiar velocities of SDSS-II SNe and found no detectable effect (see Appendix B). We therefore do not include peculiar velocity correlations in further analyses. Besides peculiar velocities we ignore further potential correlations between individual SNe and treat them according to Eq. (1) as statistical independent.

Following standard practice, we add an intrinsic dispersion in distance modulus, σ_{int}^2 , in the denominator of Eq. (1) and determine it by setting $\chi^2/N_{\text{dof}} = 1$ for the best-fit cosmological model. This term accounts for the fact that the errors on the distance moduli reported by the light curve fitter could underestimate the real error if we assume that a smooth cosmology is the correct underlying model. For the best-fit model with constant w , we find $\sigma_{\text{int}} = 0.088 \text{ mag}$ for our SDSS-only SN sample, while in Kessler et al. (2009) a value of $\sigma_{\text{int}} = 0.16$ derived from the nearby supernova samples is used in combination with the SDSS data. This results in broader contours compared to the one shown in this paper (see their Fig. 26a, and Appendix E for a possible explanation), but only marginally changes the most likely values of w and Ω_M . Similarly if we omit any intrinsic dispersion we get narrower contours but only slight changes in w and Ω_M well within the errors given in Table 1.

A further complication arises due to the uncertain calibration of the absolute magnitude (at peak) of SNe Ia, leading to a degeneracy with the absolute value of H_0 . To account for this, we marginalize over H_0 . This procedure makes use of the relative distances reported by the light curve fitter but not of their absolute value. In a recent paper Riess et al. (2009) redetermined the Hubble constant H_0 or equivalently the absolute brightness of SN Ia. One could use this value obtained from measurements in the nearby universe as a prior on H_0 in our analysis. But - as laid out in the introduction - we want to limit our analysis to probes taken at comparable redshifts as the SDSS SNe and therefore refrain from using this additional information.

3 TESTING COSMIC ACCELERATION

Given the homogeneity of the SDSS-II SN dataset, we begin our analysis by revisiting the original evidence for cosmic acceleration in the expansion of the universe. In detail, if we simply assume the universe is homogeneous and isotropic, and is described by a Robertson-Walker metric with a scale factor of $a(t)$, then the purely kinematic deceleration parameter, q , is defined by

$$q(z) \equiv -\frac{\ddot{a}a}{\dot{a}^2}, \quad (3)$$

where $q < 0$ corresponds to acceleration (we also assume light follows null geodesics and is therefore redshifted in the usual way). We relate q to the Hubble parameter by

$$H(z) \equiv \frac{\dot{a}}{a} = H_0 \exp\left(\int_0^z \frac{1+q(z')}{1+z'} dz'\right). \quad (4)$$

Luminosity distance can then be calculated directly from the expansion history via

$$d_L(z) = c(1+z) \int_0^z \frac{dz'}{H(z')}, \quad (5)$$

where we have assumed a flat universe. Thus, the magnitudes and redshifts of any SNe can be used to constrain $q(z)$ without choosing a particular dark-energy model, or even a particular theory of gravity. The assumption of flatness is necessary in practice since constraints on $q(z)$ degrade significantly when curvature is allowed to vary. A prior on curvature from CMB measurements would, of course, strengthen the constraint, but such a prior is based on the validity of GR, counter to the intention of the $q(z)$ analysis.

The simplest deceleration model we can fit is a constant, $q(z) = q_0$. In this case, the luminosity distance simplifies to

$$d_L(z) = \frac{c(1+z)}{H_0 q_0} \left[1 - (1+z)^{-q_0}\right]. \quad (6)$$

Fitting Eq. (6) to the SDSS-only SN data, we find a best fit of $q_0 = -0.34$ and $h = 0.636$. Marginalizing the joint probability density function (PDF) over h , we find $q_0 = -0.34 \pm 0.18$, or $q_0 < 0$ with 97% probability, i.e., the SDSS alone finds evidence for acceleration at 2σ without concerns regarding the absolute calibration of the peak brightness of SNIa and the relative calibration between SN surveys.

There is no reason to expect $q(z)$ to be constant; indeed, under Λ CDM, q evolves from $1/2$ during matter domination to -1 when vacuum energy dominates. However, we find that including additional parameters in $q(z)$ (e.g. by Taylor expanding q) degrades our constraints to the point of being uninteresting.

3.1 Principal Component Analysis

When trying to reconstruct an unknown function from noisy data, there is the concern that particular features of the reconstruction are not indicative of the true underlying function, but an artifact of the chosen parameterization. This concern is magnified for a function like $q(z)$, which is related to the data $\mu(z)$ by two derivatives. Therefore, we need a more robust way to determine if the universe has accelerated — a way that does not depend on the naive assumption of a constant $q(z)$. A principal-component analysis

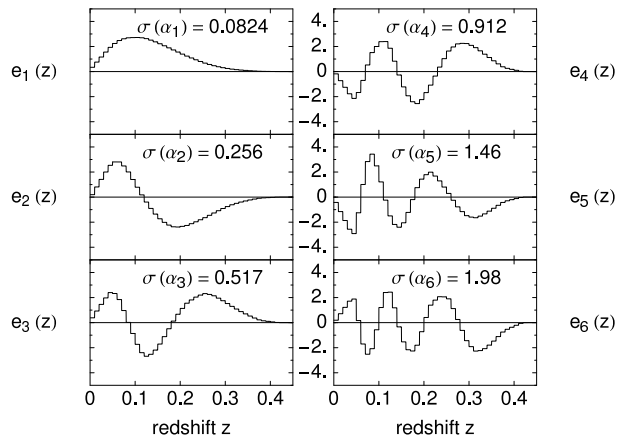


Figure 2. The principal components for $q(z)$ derived from the SDSS-II SN data. Each mode includes a Fisher matrix estimation of the error bar, $\sigma(\alpha_i)$, for its coefficient α_i . The errors are uncorrelated, and we have ordered the modes according to the size of their error bars. Only the first six modes are shown.

(Huterer & Starkman 2003) can be used to address this issue in a parameter-independent way. Principal components are a unique set of orthogonal basis functions ($e_i(z)$), such that

$$q(z) = \sum_i \alpha_i e_i(z), \quad (7)$$

which allows us to specify $q(z)$ using the coefficients, α_i . The principal components, or “modes” $e_i(z)$, are explicitly constructed so that the coefficients can be measured independently of each other, i.e., they have uncorrelated error bars. To construct these modes, we start with a piecewise-constant parameterization of $q(z)$ in bins of $dz = 0.01$, and we use our data to calculate a Fisher matrix for this parameter set and the Hubble parameter H_0 . After marginalizing over H_0 , the modes $e_i(z)$ are given by the eigenvectors of the resulting matrix. We are free to normalize these functions so that

$$\int e_i(z) e_j(z) dz = \delta_i^j, \quad (8)$$

which now specifies each function up to an overall sign convention. This procedure is completely general; in the limit of $dz \rightarrow 0$, we can specify any continuous function using these modes. The procedure is parameter-independent in the sense that we do not specify the modes a priori: they are determined primarily by the data. We do choose $q(z) = 0$ as the fiducial model for our Fisher matrix calculation, but the resulting modes are insensitive to this choice.

The six modes constrained best by the SDSS-only SN data are shown in Fig. 2. Each mode is accompanied by an estimate of the error bar on its coefficient, $\sigma(\alpha_i)$, calculated from the eigenvalues of our Fisher matrix. The first coefficient, α_1 , has the tightest constraint, while higher mode coefficients have increasingly larger errors and so provide little information about the shape of $q(z)$. Note that the first mode does not cross zero, and we have chosen its sign to be wholly positive. As argued by Shapiro & Turner (2006), this mode is useful since in addition to being well constrained, α_1 can only be negative if $q(z)$ is negative for some z . There-

fore, if we fit the model $q(z) = \alpha_1 e_1(z)$ to the data, and then measure $\alpha_1 < 0$, it constitutes parameter-independent evidence that the universe has accelerated at some redshift. We need not consider additional modes since the constraint on α_1 is, by construction, independent of the constraints on the other α_i . Marginalizing over h , we find $\alpha_1 = -0.155 \pm 0.086$ and determine that $\alpha_1 < 0$ with 96% probability. The fact that the error bar from our fit closely matches the Fisher matrix estimate of $\sigma(\alpha_1) = 0.082$ demonstrates consistency. This result is comparable to our q_0 fit above, while providing robust evidence for an accelerating universe regardless of the model assumptions.

The significance of our detection could be enhanced by combining the SDSS-II SNe with other SN datasets. This has already been done in part by Shapiro & Turner (2006) but could suffer from systematic uncertainties associated with combining data from different instruments and surveys. Again, the reader is referred to Kessler et al. (2009) for a detailed description of such combinations of the SDSS-II SN Survey with other SN datasets. We plan to repeat this analysis with the full 3-year SDSS-II SN dataset.

4 COMPARISON OF DISTANCES

We next consider the comparison of our SDSS-II SN distances with other geometrical distance estimates over the same redshift range. This is motivated by the original results of Percival et al. (2007) who noticed some tension (at $> 2\sigma$) between the cosmological constraints derived from nearby BAO measurements ($z = 0.2$ and $z = 0.35$) and higher redshift SNe of Astier et al. (2006). The BAO provide a measure of distances in the universe by relating the scale of the sound horizon at last scattering (r_s) to the scale of the corresponding correlations seen in the galaxy distribution. One such measurement of this ratio is given by the A -parameter in Eisenstein et al. (2005). This parameter is frequently used in combination with SN data to derive constraints on w (Astier et al. 2006; Kowalski et al. 2008), e.g., see Kessler et al. (2009) for the combination of the SDSS-II SN data with the A -parameter.

Here, we adopt the quantity

$$D_V(z) = \left[D_M^2 \frac{cz}{H(z)} \right]^{\frac{1}{3}}, \quad (9)$$

defined in Eisenstein et al. (2005) where D_M is the co-moving distance. Percival et al. (2007) showed by combining measurements of r_s/D_V at redshifts of 0.2 and 0.35 from both the 2dFGRS and SDSS galaxy samples, that one can obtain the ratio of the distance between two different redshifts that is independent of both r_s and H_0 . This approach also avoids a large extrapolation between the redshift of recombination ($z_{\text{CMB}} = 1090$, Komatsu et al. 2008) and these intermediate-redshift measurements.

For the analysis presented in this paper, we have adopted the latest value of $D_V(z = 0.35)/D_V(z = 0.2) = 1.736 \pm 0.065$ taken from Percival et al. (2009). The inferred value of $D_V(z = 0.35)/D_V(0.2)$ from Percival et al. (2009) is lower than that of Percival et al. (2007), bringing it into better agreement with Λ CDM. This change was caused by a revised error analysis and a change in the methodology

adopted, as well as the addition of more data. In this paper we do not use constraints on $r_s(z_d)/D_V(z)$, which depend on the sound horizon at the baryon drag epoch $r_s(z_d)$. We therefore avoid including CMB data, commonly used to model the sound horizon.

4.1 The Sliding Window Method

In Fig. 1 we show the Hubble diagram for the SDSS-II SN data discussed in Section 2 compared with a variety of cosmological and non-parametric models discussed herein. We find a scatter of 0.14 mag around these models independent of the particular fitting method. The first non-parametric model we consider is a “sliding window” method which allows us to investigate the general shape and smoothness of the Hubble diagram without assuming a cosmological model. We have thus fit piecewise Hubble parameters and luminosity distances in different redshift bins using a local redshift window following the approach of Daly & Djorgovski (2003). At each redshift z , we fit the SNe co-moving distances, $D_M(z_i) = d_L/(1+z_i)$, over a fitting window of $z_i - \Delta z < z_i < z_i + \Delta z$ ¹, by a polynomial of second order given by:

$$f(z_i - z) = A_0 + A_1(z_i - z) + A_2(z_i - z)^2. \quad (10)$$

The values of A_i are determined separately via χ^2 -minimisation in each redshift window, and we slide this window as a function of redshift in increments of 0.01 throughout the entire range. The best-fit D_M at z is proportional to A_0 , while the best-fit c/H is proportional to A_1 , and D_V is related to $[zA_0^2A_1]^{1/3}$. Our results depend on the size of the redshift window, with a wider window allowing less flexibility but smaller errors, and vice versa.

We show in Fig.1 the resulting non-parametric fit to the SDSS-II SN data as an example for a window size of $\Delta z = 0.15$ which demonstrates that the SDSS SNe data is fully consistent with the individual BAO measurements at $z = 0.2$ and 0.35 of Percival et al. (2009) (the cyan and green shaded regions indicate the 1σ and 2σ errors, which are highly correlated as they share the same data points in the overlapping fitting windows). Next, we derive the ratio $D_V(0.35)/D_V(0.2)$ and determine the covariances between the A_i values within, and between, redshift bins using the observational errors. This is shown in Fig. 3 for different window sizes. It is interesting to note that the sliding window method tends to prefer large values (steeper slopes) for D_V compared to D_V calculated from a q_0 fit described in Section 3, or D_V from the best fitting $w - \Omega_m$ parametrizations given Table 1. We also see that the sliding window method provides values for D_V which are fully consistent with the BAO result. We also show in Fig. 3 the best-fit ratios of D_V when the SDSS-II SN data are simultaneously fit with the BAO data. The SN constraints dominate these results because of their smaller uncertainty.

¹ We actually allow for a tapered region at both ends of this redshift window where the weights (i.e., the inverse-squared errors) are reduced by a Gaussian function with a standard deviation of $\sigma_z = 0.02$. This leads to a suppression of fluctuations due to the inclusion of individual data points into the window (see Daly & Djorgovski 2003).

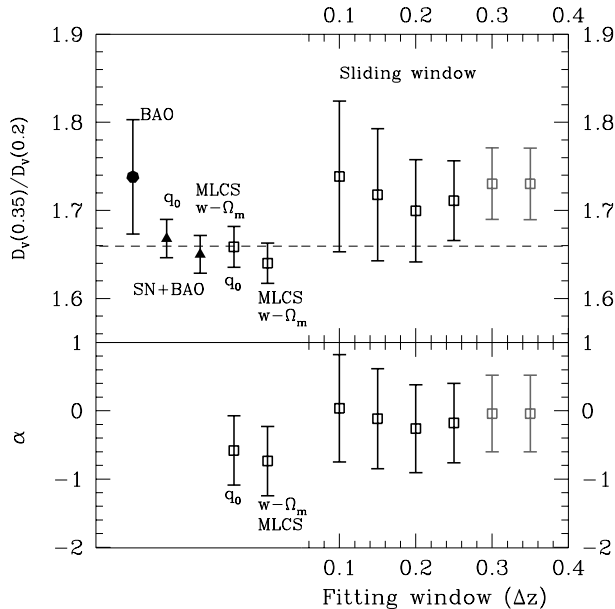


Figure 3. Upper panel: Measurements of the ratio of $D_V(0.35)/D_V(0.2)$ using different fits to the SN data, in comparison to or in combination with the BAO (solid circle) measurement. The points labelled q_0 indicate D_V derived using the fit in Section 3. Points labelled $w - \Omega_M$ are derived from the $w - \Omega_M$ parameterization marginalized over $0 < \Omega_M < 1$ and $-2 < w < 0$. Points shown as triangles include the BAO measurement as a prior whereas boxes are without the BAO. The points to the right show the values derived from the “sliding window” method as a function of the window width. The grey points are for values of the sliding window that are comparable in size to the redshift range of the whole SDSS data-set and therefore, approach the global fit. The dashed line indicates Λ CDM with $\Omega_M = 0.3$ and $\Omega_\Lambda = 0.7$. Lower panel: Given the above-mentioned parameterizations, we show the best-fit value of α defined in Eq. (13).

4.2 Testing the Distance Duality Relation

In the following we use the famous reciprocity relation (Etherington 1933; Ellis 1971), or distance duality to compare the SN and BAO distance scales. In detail, the angular diameter distance and luminosity distance are related by

$$\frac{d_L}{d_A} = (1+z)^2 \quad (11)$$

(for a discussion, see e.g. Bassett & Kunz 2004). This relation relies on photon conservation, but holds for any geometry and any metric theory of gravity where photons follow null geodesics. Therefore, it is a general test of our underlying assumptions about the nature of our Universe.

One might have expected that the distance duality relation has already been tightly constrained by observations of the blackbody CMB spectrum from the COBE FIRAS experiment (Mather et al. 1994). However, this observation does not constrain deviations from distance duality as the photon number may not be conserved (either through production or loss of photons) or more radically, photons may not follow null geodesics. Also, a grey dust component that absorbed photons independent of frequency would not cause spectral distortions away from a blackbody in the CMB since all frequencies would be affected equally. However, this grey dust would cause strong devi-

ations from distance duality since it would make the luminosity distance to any objects larger while leaving the angular diameter distance unchanged. Another way to hide the distance duality effects from CMB observations would be to affect photon number only at much higher or lower frequencies than the microwave. This is, for example, what was needed to make the axion-photon mixing proposal for the dimming of the SNe Ia consistent with CMB constraints (Csáki, Kaloper, & Terning 2002).

Beyond the CMB, several other analyses, using similar data to that discussed herein, have reported evidence for violations of distance duality at the $\sim 2\sigma$ level (Bassett & Kunz 2004; Lazkoz, Nesseris & Perivolaropoulos 2008). We revisit this issue here using a methodology similar to that outlined by More, Bovy, & Hogg (2009) and Avgoustidis, Verde, & Jimenez (2009) that does not rely on the absolute calibration of the distances to compute the ratio $d_L(z)/(d_A(z)(1+z)^2)$. Instead, we check the relative behaviour of this ratio as a function of redshift by testing the consistency of the ratio at two redshifts, $z = 0.2$ and $z = 0.35$ where we now have updated BAO measurements from Percival et al. (2009).

In the following we parameterize the distance duality relation in what we call the α -model as

$$d_L = (1+z)^{2+2\alpha} d_A = (1+z)^{(1+\alpha)} D_M, \quad (12)$$

where $\alpha = 0$ represents the expected distance duality relation, and therefore $\alpha \neq 0$ indicates a possible violation. To quantify the discrepancy between the two measures, we replace D_M in Eq. (9) with d_L from Eq. (12), and derive the relation

$$\frac{(1+z_2)^{2+2\alpha} z_1}{(1+z_1)^{2+2\alpha} z_2} = \frac{D_V^3(z_1) d_L^2(z_2) H(z_1)}{D_V^3(z_2) d_L^2(z_1) H(z_2)}. \quad (13)$$

On the right-hand side, $D_V(z_1)/D_V(z_2)$ is given by the BAO measurements, while $d_L^2(z_2)/d_L^2(z_1)$ and $H(z_1)/H(z_2)$ have to be inferred from the SDSS-II SN data. We calculate d_L and $H(z)$ using the two methods introduced above. First, we use the results from the q_0 fitting, and subsequently we calculate $d_L(z, q_0)$ and $H(z, q_0)$ at redshifts of 0.2 and 0.35. The parameter α , and its error, are then calculated by Eq. (13). We note that Eq. 13 probes the consistency of the ratio at one redshift given the ratio at the other redshift and is therefore not sensitive to any scaling proportional to $(1+z)^2$ but would be sensitive to any other loss function.

For the second method, we use the sliding window technique to derive $\vec{D}_{V, \text{SN}} \equiv [zA_0^2 A_1]^{1/3}$ (see Eq. (10)) and the corresponding covariance matrix at the two redshifts ($z = 0.2$ and $z = 0.35$) where we have BAO measurements. The best fit and error of α is calculated by applying Bayes theorem. In detail we model $\vec{D}_{V, \text{BAO}} = \beta(1+z)^{2\alpha/3} \vec{D}_{V, \text{SN}}$ at the two redshifts based on Eq. 13 where β is a free scale parameter absorbing H_0 and the scale of the sound horizon r_s at recombination. We then calculate the likelihood of the BAO \vec{D}_V measurements for α and β , by integrating over all possible SN \vec{D}_V at $z = 0.2$ and $z = 0.35$ given the gaussian prior $p(\vec{D}'_V | \vec{D}_{V, \text{SN}})$ constructed from the results of the sliding window technique:

$$L(\vec{D}_{V, \text{BAO}} | \beta, \alpha) = \int d\vec{D}'_V \times$$

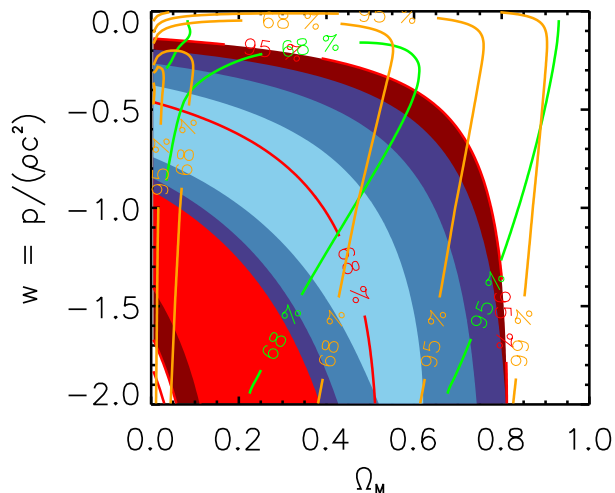


Figure 4. Confidence intervals (68%, 95%, and 99%) in the $w - \Omega_M$ plane for a flat universe and a w CDM model derived from the SDSS SNe (shaded blue region), the BAO $D_V(0.35)/D_V(0.2)$ ratio (filled red, upper 99% contours are off the plot), redshift-space distortions (orange), and the ISW effect (green). In regions where the SN contours overlap the BAO contours the latter are indicated as red lines.

$$\left[L(\vec{D}_{V, \text{BAO}} | \beta(1+z)^{2\alpha/3} \vec{D}'_V) p(\vec{D}'_V | \vec{D}_{V, \text{SN}}) \right], \quad (14)$$

The covariance matrix used in $L(\vec{D}_{V, \text{BAO}} | \beta(1+z)^{2\alpha/3} \vec{D}'_V)$ is taken from Percival et al. (2009). Applying this procedure and subsequently marginalizing over β , we calculate the best fit α and its error. The results are shown in the lower panel of Fig. 3, along with the result from the q_0 method.

For the q_0 parameterization, we find $\alpha = -0.55 \pm 0.45$, while for the sliding window scheme, we find that all results agree with $\alpha = 0$ within one sigma, i.e., the errors on α are $\simeq 0.5$ for most of the window sizes shown in Fig. 3. These results re-enforce our findings that the SN and BAO distance scales are now in good agreement over this redshift range (confirming the new findings of Percival et al. (2009)).

5 CONSTRAINING COSMOLOGICAL PARAMETERS

In contrast to the previous sections, which focused on kinematic constraints of the cosmic expansion, here we investigate the constraints on standard cosmological parameters using the SDSS SN data only in combination with dynamical measurements, eg. from the BAO, redshift space distortions and ISW. This will have less statistical power than using the combination of SN datasets presented in Kessler et al. (2009) and Sollerman et al. (2009) but our analysis is complementary to these companion papers and maximizes the impact of the homogeneity of the SDSS data (both for the SN and BAO).

In Fig. 4, we begin by showing our constraints on w and Ω_M using only the SDSS SN data. This is achieved using the χ^2 statistic according to Eq. (1) over a three dimensional grid of 200 bins in w , 100 bins in Ω_M and 160 bins in H_0 ranging from 40 to 80 $\text{km s}^{-1} \text{Mpc}^{-1}$. Subsequently, we convert the χ^2 into a likelihood using $L = \exp(-\frac{1}{2}(\chi^2 - \chi_{\text{min}}^2))$, where χ_{min}^2 is the minimum χ^2 found in the parameter space

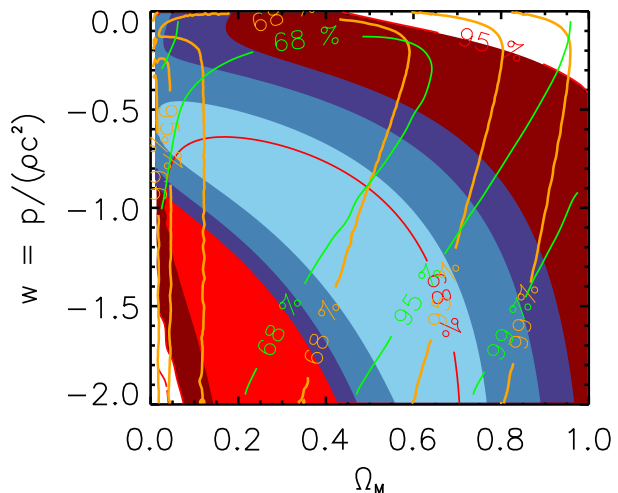


Figure 5. Similar to Fig. 4, but now including curvature as described in Appendix A.

and, in our case, is by definition close to the number of SNe in the dataset (as we have added σ_{int}). We then marginalize over H_0 by summing the likelihoods over the H_0 -bins. The shaded blue contours show the resulting confidence levels (68%, 95%, and 99%) in the $w - \Omega_M$ plane under the assumption of a flat universe ($\Omega_\Lambda = 1 - \Omega_M$). These contours include statistical errors only; systematic errors are discussed in Section 5.2.

Fig. 5 is similar, but here - instead of assuming a flat universe - we allow for curvature according to the CMB shift parameter R . We do this by calculating for a given (w, Ω_M) combination the corresponding value of Ω_Λ according to the constraints from R (see Appendix A for more details on the treatment of curvature). As discussed in Section 1, we see little effect on our results from allowing curvature to vary because of the relative small deviations from flatness allowed by the WMAP data. For results see Table 1.

For comparison with the blue SN contours, we also provide in Fig. 4 and Fig. 5 the (red) contours for the BAO measurements from Percival et al. (2009). The BAO measurements are in reasonable agreement with the SDSS-only SN contours but still prefer $w < -1$ as originally discussed in Percival et al. (2007), although all measurements are consistent with a cosmological constant. As the BAO and SN distances have reasonable overlap, we provide in Table 1 various constraints on w and Ω_M derived from combinations of the SN data with intermediate redshift dynamical measurements described in the following Sections and the BAO. We note that the SN data dominates the width and position in the w -direction of the likelihood contours, as they have smaller errors compared to the BAO constraints, while adding the BAO data helps to curtail the large values of Ω_M seen for the SDSS-only results.

5.1 Constraints from dynamical measurements

To improve the cosmological constraints, we can include a number of other low redshift cosmological measurements thus providing a first measurement of the cosmological model within the local universe. In particular, we consider constraints on w derived from measurements of the growth

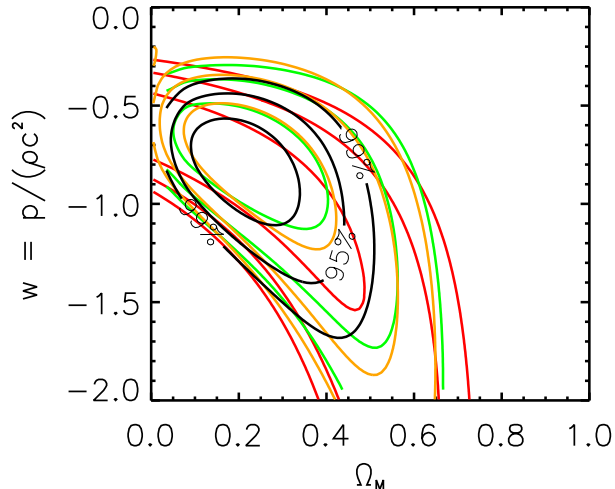


Figure 6. Confidence intervals (68%, 95%, and 99%) in the $w - \Omega_M$ plane for the combination of SDSS-II SNe with BAO (red), redshift-space distortions (orange), ISW effect (green) and for the combination of SN, BAO, RS and ISW (black) under the assumption of a flat universe. Numerical results are given in Table 1.

of structure in the universe, including redshift-space distortions from the 2dFGRS and the ISW effect from the SDSS imaging survey. Both of these methods are particularly sensitive to Ω_M and thus provide independent and orthogonal constraints to the SN data.

5.1.1 Redshift-Space Distortions

Under the assumption that galaxies are related to the large scale dark-matter distribution, the anisotropy of the redshift-space correlation function depends on the parameter

$$\beta(z) = f_g(z)/b(z), \quad (15)$$

where $b(z)$ is the linear bias relating the galaxies to the underlying dark matter and $f_g(z)$ the growth rate of structure. Absolute deviations between the real-space and redshift-space correlation functions depend on the parameter combination $f_g(z)\sigma_8(z)$, where $\sigma_8(z)$ is defined as the root-mean square (rms) mass fluctuation in spheres of radius $8h^{-1}$ Mpc, and provides a convenient way of normalising the matter fluctuations (for a recent review see Percival & White 2008).

To remove either the dependence on galaxy bias from a measurement of $\beta(z)$, or equivalently the dependence on $\sigma_8(z)$ from the measurement of $f_g(z)\sigma_8(z)$, we need further cosmological information. In this paper, we adopt the central value of $\beta(z \approx 0.15) = 0.49 \pm 0.09$ calculated by Hawkins et al. (2003) from the 2dFGRS, which is an update of the measurement given by Peacock et al. (2001). We follow the procedure outlined by Guzzo et al. (2008), and convert from β to f_g by adding an additional uncertainty of 0.12 in quadrature to account for the uncertainty in galaxy bias, which is estimated to be close to unity. This error includes the cosmological dependence of the bias measurement. For further calculations, we thus use $f_g(z = 0.15) = 0.49 \pm 0.15$,

Table 1. Results for w and Ω_M from combinations of the SDSS-II SN and BAO data with redshift-space distortions (RS) and the ISW effect. We provide measurements based on the MLCS2k2 light-curve fitting method and our assumptions about curvature (Geo.). Values for w are derived after marginalizing over Ω_M and vice versa. The given uncertainties are statistical errors only; for an estimate of the systematic uncertainty see Section 5.2.

Dataset	Geo.	w	Ω_M
SN+BAO	flat	$-0.74^{+0.17}_{-0.32}$	$0.37^{+0.16}_{-0.64}$
SN+RS	flat	$-0.77^{+0.19}_{-0.25}$	$0.26^{+0.15}_{-0.12}$
SN+ISW	flat	$-0.74^{+0.16}_{-0.22}$	$0.23^{+0.15}_{-0.12}$
SN+RS+ISW	flat	$-0.76^{+0.17}_{-0.19}$	$0.23^{+0.10}_{-0.08}$
SN+RS+ISW+BAO	flat	$-0.81^{+0.16}_{-0.18}$	$0.22^{+0.09}_{-0.08}$
SN+BAO	curved	$-0.99^{+0.30}_{-0.59}$	$0.50^{+0.13}_{-0.22}$
SN+RS	curved	$-0.82^{+0.19}_{-0.26}$	$0.31^{+0.15}_{-0.12}$
SN+ISW	curved	$-0.78^{+0.16}_{-0.22}$	$0.27^{+0.15}_{-0.13}$
SN+RS+ISW	curved	$-0.79^{+0.16}_{-0.20}$	$0.27^{+0.10}_{-0.09}$
SN+RS+ISW+BAO	curved	$-0.85^{+0.17}_{-0.19}$	$0.27^{+0.10}_{-0.09}$

assuming the weighted median redshift of the 2dFGRS survey of $z \approx 0.15$.

We adopt the parameterization given by Linder (2005),

$$g(a) = \frac{\delta}{a} = \exp \int_0^a d \ln a [\Omega_M(a)^\gamma - 1], \quad (16)$$

which is related to the growth rate by

$$f_g = \frac{\dot{\delta}}{g}, \quad (17)$$

where $\delta = \delta\rho_M/\rho_M$ describes the perturbations in the density of matter (ρ_m). For constant $w > -1$, the exponent γ in Eq. (16) can be approximated as

$$\gamma = 0.55 + 0.05(1 + w), \quad (18)$$

while for a phantom-like dark energy component (with $w < -1$) the exponent is

$$\gamma = 0.55 + 0.02(1 + w). \quad (19)$$

Solving Eq. (17) numerically, we derive the (orange) contours shown in Fig. 4 and Fig. 5. In Table 1, we present the constraints on w from the combinations of these data with the SDSS-II SN and BAO likelihoods, marginalized over Ω_M and vice versa.

5.1.2 The Late-Time integrated Sachs-Wolfe Effect

The integrated Sachs-Wolfe (ISW) effect is caused by the change in energy of CMB photons as they pass through a time-varying gravitational potential (Sachs & Wolfe 1967). In a flat, matter-dominated universe, we would not expect to see an ISW effect as the large-scale gravitational potentials do not change in conformal time. However, in a universe dominated by DE or curvature, we should detect a so called *late-time* ISW effect, which provides a direct measure of these quantities at the redshift of the changing potentials, i.e., the effect does not depend on the previous history of the growth of structure.

The late-time ISW effect introduces additional secondary anisotropies on top of the primary CMB fluctuations

and is therefore hard to detect directly. However, the ISW effect can be seen via the cross-correlation of the CMB with tracers in the large-scale structure of the universe as outlined by Crittenden & Turok (1996). This has now been achieved by a number of authors using a host of different galaxy datasets (Fosalba, Gaztañaga, & Castander 2003; Scranton et al. 2003; Boughn & Crittenden 2004; Nolta et al. 2004; Afshordi, Loh, & Strauss 2004; Fosalba & Gaztañaga 2004; Padmanabhan et al. 2005; Cabré et al. 2006; Giannantonio et al. 2006; Rassat et al. 2007; Ho et al. 2008; Granett, Neyrinck, & Szapudi 2008).

In this paper, we exploit the recent analysis of Giannantonio et al. (2008) and focus on the subset of intermediate-redshift ($z < 0.4$) SDSS data they used. Even this subset of data shows a detection of the ISW effect at the 3σ level (Giannantonio 2008a). The contours for this new determination of the ISW effect are plotted in Figs. 4, 5, and 6. In Table 1, we present the combination of this new ISW effect measurement with our SDSS-II SN and BAO data, using the same procedure as discussed in Section 5.1.1.

5.2 Evaluation of SN Systematics

We provide in Table 1 measurements of w and Ω_M from various combinations of the four data-sets considered herein (SN, BAO, RS, ISW). The most stringent constraint comes from the combination of all the probes giving $w = -0.81^{+0.16}_{-0.18}$ and $\Omega_M = 0.22^{+0.09}_{-0.08}$, which is competitive given the restricted redshift range considered in this analysis. However, much of this constraint comes from the combination of just the SDSS-II SNe and ISW measurements (See Table 1). The ISW contours already include correlations between different angular and redshift bins and cosmic variance, and could therefore be considered stable (see e.g. Giannantonio et al. 2008). Similarly the contours we use for redshift-space distortions include the dominant uncertainty coming from the galaxy bias (see the procedure laid out in Section 5.1.1). Percival et al. (2009) has done several checks and found that their result is robust against variations in sample selection, number of redshift slices, calibration and other potential influences. Therefore, the results presented in Table 1 includes major uncertainties affecting the other probes but only the statistical uncertainties from the SDSS-SN data on the measured cosmological parameters.

As discussed in Kessler et al. (2009), the SDSS-II SN distances also depend on the detailed choices and assumptions within the MLCS2k2 supernova light-curve fitting procedure, including different training vectors, priors on A_V and R_V , uncertainties in zero points and the filter systems, and selection biases. To quantify the systematic uncertainties associated with these parameter choices, we repeat our analysis above for these different choices and, following the procedure laid out in Kessler et al. (2009), we calculate a variation of $\Delta w = \pm 0.15$ with respect to the fiducial MLCS2k2 reduction presented in Table 1 in the case of combining all four constraints and slightly larger values for the other cases.

Our estimates of the systematic uncertainty for MLCS2k2 are larger than the values calculated in Kessler et al. (2009) because they use the BAO A -parameter from Eisenstein et al. (2005), and the added constraints on Ω_M and w derived from using the CMB R -parameter (Komatsu et al. 2008). We reproduce their values for the

systematic uncertainties on MLCS2k2 ($\Delta w \approx 0.1$), if we include these constraints in our analysis. However, in this paper, we restrict our analysis to intermediate-redshift probes and therefore do not include the CMB constraints, which results in larger uncertainties.

Our analysis of the MLCS2k2 systematic uncertainties discussed above does not include the large shift in w discussed in Kessler et al. (2009) when the rest-frame U -band template is removed in the light-curve fitting. As seen in Table 6 of Kessler et al. (2009), removing the rest-frame U -band results in a -0.31 shift in w , while we find a shift of -0.43 if we remove this data. This particular uncertainty would therefore give rise to a bimodal result either centered around $w \simeq -0.8$ (with U -band included) or $w = -1.2$ (without U -band), yet both consistent with $w = -1$ within the errors.

We do not add the uncertainty due to excluding the rest-frame U -band to our systematic errors because we believe it is incorrect to exclude this data from the light-curve fitting. Even though there is evidence for diversity in the UV spectra of SNe Ia (see Ellis et al. 2008; Foley et al. 2008), the removal of the rest-frame U -band data from the SDSS-only analysis results in the light curve fitter using only two filters at $z < 0.2$ to constrain the colors of the SNe. This provides significant freedom to the MLCS2k2 fitter and then the priors on the fitted parameters become important. We note that w is only shifted by -0.1 when using the SALT light curve fitter (see Table 8 in Kessler et al. 2009) on the SDSS-only sample with the rest-frame U -band excluded. This is the only noticeable difference between these two light curve fitting methodologies when considering the SDSS-only sample; namely the error on w when the rest-frame U -band is excluded. Finally, Kessler et al. (2009) also sees a clear jump in the SDSS Hubble diagram at $z \simeq 0.2$ when the rest-frame U -band is excluded from the MLCS2k2 analysis (see their Section 10.1.3 and Fig. 30), indicating that a constant w model is not a good fit in this case.

6 CONCLUSIONS

We present an analysis of the luminosity distances of Type Ia Supernovae from the Sloan Digital Sky Survey-II (SDSS-II) Supernova Survey in conjunction with other intermediate redshift ($z < 0.4$) cosmological measurements including redshift-space distortions from the 2dFGRS, the ISW effect, and the BAO distance scale from both the SDSS and 2dFGRS. We have analyzed the SDSS-II SN luminosity distances using several 'model-independent' methods, including fitting the data using a $q(z)$ parameterization, principal components, and a non-parametric "sliding window" method. We find consistent results between all these methods that provides evidence for an accelerating universe based solely on the first-year SDSS-II SN data. The strongest evidence we find comes when we make the strongest assumptions, that q_0 is constant and the universe is flat which gives probability for acceleration of $> 97\%$.

We also compare our SDSS-II SN data with the local BAO measurements, and find they are in good agreement. This is in contrast with the findings of Percival et al. (2007) which found tension between the two distance measures, but confirms the new BAO analysis of Percival et al. (2009) who

note that this tension has now lessened. Taking this observation further, we test the distance duality relation, i.e., for any metric theory of gravity, we expect $d_L/(d_A(1+z)^2) = 1$. We see no evidence for a discrepancy from this relation (at the one sigma level) in contrast to previous claims for a potential violation on the 2σ level as seen in (Bassett & Kunz 2004; Lazkoz, Nesseris & Perivolaropoulos 2008). Finally, we present a new measurement of the equation-of-state parameter of dark energy using a combination of geometrical distances in the universe and estimates for the growth rate of structure. Our strongest constraint comes from the combination of all four data-sets discussed herein (SDSS-II SN, BAO, redshift-space distortions, ISW) with $w = -0.81_{-0.18}^{+0.16}$ (stat) and $\Omega_M = 0.22_{-0.08}^{+0.09}$ (stat) (assuming a flat universe). However, the combination of just the SDSS-II SNe and the ISW measurements alone is almost as powerful in constraining these parameters (Table 1). Our results only change slightly if we allow curvature to vary, consistent with the CMB measurements (see Appendix A). We quote a systematic uncertainty of $\Delta w = \pm 0.15$ based on the details of the MLCS2k2 light-curve fitter (see Kessler et al. 2009 for a fuller discussion).

Thus we have shown that low-redshift cosmological probes give a self-consistent picture of the distance-redshift relation. When combined with growth of structure and ISW at the same epoch that picture is consistent with Λ CDM and re-enforces the complementarity amongst other data and analyses in the literature.

ACKNOWLEDGEMENTS

We thank an anonymous referee for helpful comments on this paper which greatly improved the content of the paper. RCN on behalf of the authors would like to thank Mike Turner for helpful discussions on the q_0 fit and Eric Aubourg for useful discussions on the distance duality relation. We also thank Rick Kessler and Mark Sullivan for extensive discussions about their work and papers. TG thanks Jussi Väliiviita for helpful suggestions. HL, CS and RCN are grateful to STFC for funding this research with rolling grants ST/F002335/1 and ST/H002774/1. H.-J. S is supported by the D.O.E at Fermilab. A.V.F. is grateful for the support of US NSF grant AST-0607485.

Funding for the creation and distribution of the SDSS and SDSS-II has been provided by the Alfred P. Sloan Foundation, the Participating Institutions, the National Science Foundation, the U.S. Department of Energy, the National Aeronautics and Space Administration, the Japanese Monbukagakusho, the Max Planck Society, and the Higher Education Funding Council for England. The SDSS Web site is <http://www.sdss.org/>.

The SDSS is managed by the Astrophysical Research Consortium for the Participating Institutions. The Participating Institutions are the American Museum of Natural History, Astrophysical Institute Potsdam, University of Basel, Cambridge University, Case Western Reserve University, University of Chicago, Drexel University, Fermilab, the Institute for Advanced Study, the Japan Participation Group, Johns Hopkins University, the Joint Institute for Nuclear Astrophysics, the Kavli Institute for Particle Astrophysics and Cosmology, the Korean Scientist Group, the

Chinese Academy of Sciences (LAMOST), Los Alamos National Laboratory, the Max-Planck-Institute for Astronomy (MPA), the Max-Planck-Institute for Astrophysics (MPIA), New Mexico State University, Ohio State University, University of Pittsburgh, University of Portsmouth, Princeton University, the United States Naval Observatory, and the University of Washington.

This work is based in part on observations made at the following telescopes. The Hobby-Eberly Telescope (HET) is a joint project of the University of Texas at Austin, the Pennsylvania State University, Stanford University, Ludwig-Maximilians-Universität München, and Georg-August-Universität Göttingen. The HET is named in honor of its principal benefactors, William P. Hobby and Robert E. Eberly. The Marcario Low-Resolution Spectrograph is named for Mike Marcario of High Lonesome Optics, who fabricated several optical elements for the instrument but died before its completion; it is a joint project of the Hobby-Eberly Telescope partnership and the Instituto de Astronomía de la Universidad Nacional Autónoma de México. The Apache Point Observatory 3.5 m telescope is owned and operated by the Astrophysical Research Consortium. We thank the observatory director, Suzanne Hawley, and site manager, Bruce Gillespie, for their support of this project. The Subaru Telescope is operated by the National Astronomical Observatory of Japan. The William Herschel Telescope is operated by the Isaac Newton Group on the island of La Palma in the Spanish Observatorio del Roque de los Muchachos of the Instituto de Astrofísica de Canarias. The W. M. Keck Observatory is operated as a scientific partnership among the California Institute of Technology, the University of California, and the National Aeronautics and Space Administration; the observatory was made possible by the generous financial support of the W. M. Keck Foundation.

REFERENCES

- Abate A., Lahav O., 2008, MNRAS, 389, L47
- Adelman-McCarthy J. K., et al., 2008, ApJS, 175, 297
- Afshordi N., Loh Y.-S., Strauss M. A., 2004, PhRvD, 69, 083524
- Astier P., et al., 2006, A&A, 447, 31
- Avgoustidis A., Verde L., Jimenez R., 2009, JCAP, 6, 12
- Bassett B. A., Kunz M., 2004, PhRvD, 69, 101305
- Bassett B. A., Kunz M., 2004, ApJ, 607, 661
- Blake C., Glazebrook K., 2003, ApJ, 594, 665
- Blake C., Collister A., Bridle S., Lahav O., 2007, MNRAS, 374, 1527
- Boughn S., Crittenden R., 2004, Natur, 427, 45
- Cabré A., Gaztañaga E., Manera M., Fosalba P., Castander F., 2006, MNRAS, 372, L23
- Csáki C., Kaloper N., Terning J., 2002, PhRvL, 88, 161302
- Clarkson, C., Cortês, M., & Bassett, B. 2007, Journal of Cosmology and Astro-Particle Physics, 8, 11
- Cole, S., et al. 2005, MNRAS, 362, 505
- Colless M., et al., 2001, MNRAS, 328, 1039
- Conley A., et al., 2008, ApJ, 681, 482
- Conley A., Carlberg R. G., Guy J., Howell D. A., Jha S., Riess A. G., Sullivan M., 2007, ApJ, 664, L13

- Copeland, E.J., Sami, M., Tsujikawa, S., 2006., *Int. J. Mod. Phys. D*15:1753-1936
- Crittenden R. G., Turok N., 1996, *PhRvL*, 76, 575
- Daly R. A., Djorgovski S. G., 2003, *ApJ*, 597, 9
- Davis T. M., et al., 2007, *ApJ*, 666, 716
- Dunkley, J., et al. 2009, *ApJS*, 180, 306
- Eisenstein D. J., et al., 2001, *AJ*, 122, 2267
- Eisenstein, D. J., et al. 2005, *ApJ*, 633, 560
- Ellis G. F. R., 1971, in *General Relativity and Cosmology*, ed R.K. Sachs (New York: Academic Press), 104
- Ellis, R. S., et al. 2008, *ApJ*, 674, 51
- Etherington I. M. H., 1933, *PMag*, 15, 761
- Filippenko A. V., 2005, in *White Dwarfs: Cosmological and Galactic Probes*, ed. E. M. Sion, S. Vennes, & H. L. Shipman (Dordrecht: Springer), 97
- Foley, R. J., et al. 2008, *ApJ*, 684, 68
- Fosalba P., Gaztañaga E., Castander F. J., 2003, *ApJ*, 597, L89
- Fosalba P., Gaztañaga E., 2004, *MNRAS*, 350, L37
- Frieman J. A., et al., 2008, *AJ*, 135, 338
- Frieman J., Turner M., Huterer D., 2008, arXiv, 803, arXiv:0803.0982
- Fukugita M., Ichikawa T., Gunn J. E., Doi M., Shimasaku K., Schneider D. P., 1996, *AJ*, 111, 1748
- Giannantonio T., et al., 2006, *PhRvD*, 74, 063520
- Giannantonio T., Scranton R., Crittenden R. G., Nichol R. C., Boughn S. P., Myers A. D., Richards G. T., 2008, *PhRvD*, 77, 123520
- Giannantonio T., private communication
- Gordon C., Land K., Slosar A., 2007, *PhRvL*, 99, 081301
- Granett B. R., Neyrinck M. C., Szapudi I., 2008, *ApJ*, 683, L99
- Gunn J. E., et al., 1998, *AJ*, 116, 3040
- Gunn J. E., et al., 2006, *AJ*, 131, 2332
- Guy J., Astier P., Nobili S., Regnault N., Pain R., 2005, *A&A*, 443, 781
- Guy J., et al., 2007, *A&A*, 466, 11
- Guzzo L., et al., 2008, *Natur*, 451, 541
- Hamuy M., Phillips M. M., Wells L. A., Maza J., 1993, *PASP*, 105, 787
- Hamuy M., Phillips M. M., Suntzeff N. B., Schommer R. A., Maza J., Aviles R., 1996, *AJ*, 112, 2391
- Hawkins E., et al., 2003, *MNRAS*, 346, 78
- Hicken M., Wood-Vasey W. M., Blondin S., Challis P., Jha S., Kelly P. L., Rest A., Kirshner R. P., 2009, arXiv, arXiv:0901.4804
- Ho S., Hirata C., Padmanabhan N., Seljak U., Bahcall N., 2008, *PhRvD*, 78, 043519
- Hogg D. W., Finkbeiner D. P., Schlegel D. J., Gunn J. E., 2001, *AJ*, 122, 2129
- Holtzman J. A., et al., 2008, *AJ*, 136, 2306
- Holwerda B. W., 2008, *MNRAS*, 386, 475
- Howell D. A., et al., 2009, *ApJ*, 691, 661
- Hu W., Dodelson S., 2002, *ARA&A*, 40, 171
- Hu W., Haiman Z., 2003, *PhRvD*, 68, 063004
- Hui L., Greene P. B., 2006, *PhRvD*, 73, 123526
- Huterer D., Starkman G., 2003, *PhRvL*, 90, 031301
- Hütsi G., 2006, *A&A*, 459, 375
- Ivezić Ž., et al., 2004, *AN*, 325, 583
- Jha S., Riess A. G., Kirshner R. P., 2007, *ApJ*, 659, 122
- Kaiser N., 1987, *MNRAS*, 227, 1
- Kessler R., et al., 2009, arXiv, arXiv:0908.4274
- Komatsu E., et al., 2008, arXiv, 803, arXiv:0803.0547
- Kowalski M., et al., 2008, arXiv, 804, arXiv:0804.4142
- Lazkoz R., Nesseris S., Perivolaropoulos L., 2008, *JCAP*, 7, 12
- Linder E. V., 2005, *PhRvD*, 72, 043529
- Lupton R., Gunn J. E., Ivezić Z., Knapp G. R., Kent S., 2001, *ASPC*, 238, 269
- Masters K. L., Springob C. M., Haynes M. P., Giovanelli R., 2006, *ApJ*, 653, 861
- Mather J. C., et al., 1994, *ApJ*, 420, 439
- Miller C. J., Nichol R. C., Batuski D. J., 2001, *ApJ*, 555, 68
- More S., Bovy J., Hogg D. W., 2009, *ApJ*, 696, 1727
- Nolta M. R., et al., 2004, *ApJ*, 608, 10
- Nolta M. R., et al., 2008, arXiv, 803, arXiv:0803.0593
- Padmanabhan N., Hirata C. M., Seljak U., Schlegel D. J., Brinkmann J., Schneider D. P., 2005, *PhRvD*, 72, 043525
- Padmanabhan N., et al., 2007, *MNRAS*, 378, 852
- Padmanabhan N., et al., 2008, *ApJ*, 674, 1217
- Peacock J. A., et al., 2001, *Nature*, 410, 169
- Peebles P. J., Ratra B., 2003, *RvMP*, 75, 559
- Percival W. J. et al. 2007, *MNRAS*, 381, 1053
- Percival W. J., White M., 2008, arXiv, arXiv:0808.0003
- Percival W. J., et al., 2009, arXiv:0907.1660
- Perlmutter S., et al., 1997, *ApJ*, 483, 565
- Perlmutter S., et al., 1999, *ApJ*, 517, 565
- Phillips M. M., 1993, *ApJ*, 413, L105
- Pier J. R., Munn J. A., Hindsley R. B., Hennessy G. S., Kent S. M., Lupton R. H., Ivezić Ž., 2003, *AJ*, 125, 1559
- Rassat A., Land K., Lahav O., Abdalla F. B., 2007, *MNRAS*, 377, 1085
- Riess A. G., Press W. H., Kirshner R. P., 1996, *ApJ*, 473, 88
- Riess A. G., et al., 1998, *AJ*, 116, 1009
- Riess A. G., et al., 2004, *ApJ*, 607, 665
- Riess A. G., et al., 2007, *ApJ*, 659, 98
- Riess A. G., et al., 2009, *ApJ*, 699, 539
- Sachs R. K., Wolfe A. M., 1967, *ApJ*, 147, 73
- Sako M., et al., 2008, *AJ*, 135, 348
- Scranton R., et al., 2003, arXiv:astro-ph/0307335
- Seo H.-J., Eisenstein D. J., 2003, *ApJ*, 598, 720
- Shapiro C., Turner M. S., 2006, *ApJ*, 649, 563
- Smith J. A., et al., 2002, *AJ*, 123, 2121
- Sollerman J., et al., 2009, *ApJ*, 703, 1374
- Stoughton C., et al., 2002, *AJ*, 123, 485
- Strauss M. A., et al., 2002, *AJ*, 124, 1810
- Sullivan M., Ellis R. S., Howell D. A., Riess A., Nugent P. E., Gal-Yam A., 2009, *ApJ*, 693, L76
- Timmes F. X., Brown E. F., Truran J. W., 2003, *ApJ*, 590, L83
- Tonry J. L., et al., 2003, *ApJ*, 594, 1
- Tucker D. L., et al., 2006, *AN*, 327, 821
- Turner M. S., Riess A. G., 2002, *ApJ*, 569, 18
- Uzan J.-P., 2006, arXiv:astro-ph/0605313
- Wang L., Goldhaber G., Aldering G., Perlmutter S., 2003, *ApJ*, 590, 944
- Wang X., Wang L., Pain R., Zhou X., Li Z., 2006, *ApJ*, 645, 488
- Wang Y., Mukherjee P., 2007, *PhRvD*, 76, 103533
- Weller, J., & Albrecht, A. 2002, *PRD*, 65, 103512
- Wood-Vasey W. M., et al., 2007, *ApJ*, 666, 694
- York D. G., et al., 2000, *AJ*, 120, 1579

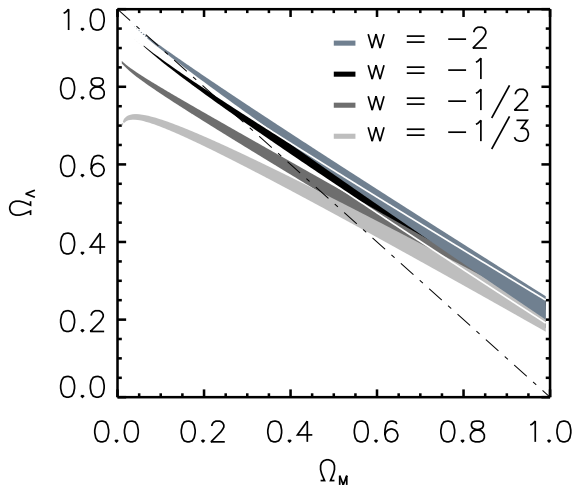


Figure A1. The 68% contours in the $\Omega_\Lambda - \Omega_M$ plane derived from the “scaled distance to recombination” R (see text) taken from the WMAP-5 results, or $R = 1.715 \pm 0.021$. The dot-dashed line indicates a flat universe.

Zheng C., et al., 2008, AJ, 135, 1766

APPENDIX A: COSMIC CURVATURE

In most parts of this paper, we assume a flat universe, or $\Omega_M + \Omega_\Lambda = 1$. Observationally, the most robust constraints on curvature come from the distance to the last scattering surface of the CMB as determined by measurements of the CMB power spectrum. In cases where we consider curvature, allowing any combination of Ω_M and Ω_Λ , we include a prior on the “scaled distance to recombination” R defined as

$$R = \sqrt{\Omega_M H_0^2} \frac{d_L(z_{\text{CMB}})}{(1 + z_{\text{CMB}})}, \quad (\text{A1})$$

given in Wang & Mukherjee (2007), with a value of $R = 1.71 \pm 0.019$ from Komatsu et al. (2008). We note that R is independent of H_0 because d_L scales linearly with $1/H_0$. For fixed values of Ω_M and w it is now possible to constrain Ω_Λ from a measurement of R . We show 68% confidence levels in Fig. A1 calculated from Eq. (A1). As R does not depend on the Hubble constant H_0 , or on the baryon density $\Omega_b h^2$, no further assumptions are needed on these quantities. From Fig. A1, it is obvious that using only R , a curvature of order a few percent cannot be neglected (depending on the value of w). We introduce the effects of curvature in the redshift-space distortion analysis by adding an additional χ^2 term in the likelihood analysis calculated from Eq. (A1) and subsequently marginalize over Ω_Λ , where as for the ISW effect we pick for a given (w, Ω_M) combination the best-fit value of Ω_Λ .

APPENDIX B: CORRELATED PECULIAR VELOCITIES

The peculiar velocities (PVs) of supernovae (SNe) introduce an additional scatter onto the Hubble diagram (see Eq. (2)). However, as pointed out by Hui & Greene (2006), we expect these PVs to be correlated, especially at low redshift, thus

leading to significant covariance between pairs of SNe, i.e., a pair of SNe at radial positions $\mathbf{r}_i, \mathbf{r}_j$ have a projected velocity correlation function of $\xi(\mathbf{r}_i, \mathbf{r}_j) = \langle (\mathbf{v}(\mathbf{r}_i) \cdot \hat{\mathbf{r}}_i)(\mathbf{v}(\mathbf{r}_j) \cdot \hat{\mathbf{r}}_j) \rangle$. We can calculate this function in linear theory using the matter power spectrum, the linear growth function and its derivative. This is interesting because, if this effect is detected, it may enable the SNe to constrain the parameters of structure formation, in addition to the standard background expansion.

The expression for the full covariance between SNe is given by Gordon, Land, & Slosar (2007); Abate & Lahav (2008)

$$C_v(i, j) = \left(1 - \frac{(1+z)^2}{Hd_L}\right)_i \left(1 - \frac{(1+z)^2}{Hd_L}\right)_j \xi(\mathbf{r}_i, \mathbf{r}_j). \quad (\text{B1})$$

This can be compared with the standard diagonal random errors, which are

$$\sigma(i)^2 = \left(\frac{\ln 10}{5}\right)^2 [\sigma_m^2 + \mu_{err}(i)^2] + \left(1 - \frac{(1+z)^2}{Hd_L}\right)_i^2 \sigma_v^2, \quad (\text{B2})$$

where the intrinsic magnitude and velocity scatters σ_m, σ_v have been introduced as usual. A numerical evaluation shows that the two are comparable at low redshift. In particular, for a pair of SNe at $z = 0.05$ and zero angular separation, the covariance is $C_v(i, j) \simeq 0.1\sigma(i)\sigma(j)$. This decreases at higher redshifts and greater separations.

This effect has been detected by Gordon, Land, & Slosar (2007) using a catalogue of 124 low redshift SNe by Jha, Riess, & Kirshner (2007) at $\bar{z} = 0.017$, and it has been carried further to constrain parameters such as σ_8 and the growth factor γ (Abate & Lahav 2008).

Here we repeated the analysis for the SDSS SNe, but since our minimum redshift is $z \simeq 0.05$, we expect the effect of correlated PVs to be small. Indeed, we found that a likelihood study performed by a Monte Carlo Markov chain (MCMC) analysis of the cosmological parameters yields no change in the results whether the full PV covariance matrix of Eq. (B1) is included or not. For example, if we set the intrinsic scatters $\sigma_m = \sigma_v = 0$, we find that the reduced χ^2/ν for the best fit cosmology decreases by only 1% when the PV covariance matrix was included.

Therefore, we are unable to detect the correlation of SN peculiar velocities with this data and are safe to ignore them in further analyses. This effect will become more important with larger samples of low redshift SNe.